

DRAFT VERSION JULY 19, 2006  
Preprint typeset using L<sup>A</sup>T<sub>E</sub>X style emulatej v. 6/22/04

# NO EXPANDING FIREBALL: RESOLVING THE RECURRENT NOVA RS OPHIUCHI WITH INFRARED INTERFEROMETRY

J. D. MONNIER<sup>1</sup>, R. K. BARRY<sup>2,12</sup>, W. A. TRAUB<sup>7,8</sup>, B. F. LANE<sup>3</sup>, R. L. AKESON<sup>6</sup>, S. RAGLAND<sup>4</sup>, P. A. SCHULLER<sup>7</sup>,  
H. LE COROLLER<sup>11</sup>, J.-P. BERGER<sup>5</sup>, R. MILLAN-GABET<sup>6</sup>, E. PEDRETTI<sup>1</sup>, F. P. SCHLOERB<sup>9</sup>, C. KORESOKO<sup>6</sup>,  
N. P. CARLETON<sup>7</sup>, M. G. LACASSE<sup>7</sup>, P. KERN<sup>5</sup>, F. MALBET<sup>5</sup>, K. PERRAUT<sup>5</sup>, M. J. KUCHNER<sup>12</sup> AND  
M. W. MUTERSPAUGH<sup>10</sup>

*Draft version July 19, 2006*

## ABSTRACT

Following the recent outburst of the recurrent nova RS Oph on 2006 Feb 12, we measured its near-infrared size using the IOTA, Keck, and PTI Interferometers at multiple epochs. The characteristic size of  $\sim 3$  milliarcseconds hardly changed over the first 60 days of the outburst, ruling out currently-popular models whereby the near-infrared emission arises from hot gas in the expanding shock. The emission was also found to be significantly asymmetric, evidenced by non-zero closure phases detected by IOTA. The physical interpretation of these data depend strongly on the adopted distance to RS Oph. Our data can be interpreted as the first direct detection of the underlying RS Oph binary, lending support to the recent “reborn red giant” models of Hachisu & Kato. However, this result hinges on an RS Oph distance of  $\lesssim 540$  pc, in strong disagreement with the widely-adopted distance of  $\sim 1.6$  kpc. At the farther distance, our observations imply instead the existence of a non-expanding, dense and ionized circumbinary gaseous disk or reservoir responsible for the bulk of the near-infrared emission. Longer-baseline infrared interferometry is uniquely suited to distinguish between these models and to ultimately determine the distance, binary orbit, and component masses for RS Oph, one of the closest-known (candidate) SNIa progenitor systems.

*Subject headings:* techniques:interferometric, stars: novae, infrared:stars, stars: individual (RS Oph)

## 1. INTRODUCTION

Most astronomers are familiar with Classical Novae, exploding stars in which an accreting white dwarf (WD) in an interacting binary system accumulates enough material for it to become unstable to hydrogen burning. The expanding blastwave from one such event Nova Aql 2005 was recently observed by infrared (IR) interferometry (Lane et al. 2005) and the geometric distance was estimated based on velocities from spectral line observations. This result is consistent with the “optically-thick fireball” model which has been successfully used for twenty years to explain the time-evolution of the spectral energy distribution of classical novae (Gehrz 1988).

While classical novae are expected to recur, very few actually have in recorded history. RS Oph is one of the handful of so-called “recurrent novae” with (now) 6 outbursts since 1898 (Warner 1976). The most recent outburst occurred on February 12, 2006 (Narumi et al. 2006), and this unusual event motivated intense monitoring by the IR interferometry community.

The special nature of RS Oph is thought to stem from two causes. First, the WD is likely extremely close to

the Chandrasekhar limit, since the amount of hydrogen needed to trigger an outburst decreases dramatically as the WD mass increases. Indeed, detailed models indicate the WD mass is within 1% of exploding as a type Ia supernova (e.g., Hachisu & Kato 2001). Second, the mass-losing companion for RS Oph is a red giant (RG) with a wind, providing a high density medium for accretion onto the WD as well as for the exploding blastwave to interact with. Bode & Kahn (1985) have produced the most successful model for recurrent novae, drawing a clear analogy to extragalactic supernovae and explaining the radio and X-ray light curves in this context.

Evans et al. (1988) were the first to study in detail the IR time evolution of a recurrent nova. They monitored closely the 1-3.5 micron flux of RS Oph for about 3 years after the 1985 eruption. They found the light curve had a characteristic (2-mag) decay time scale of about 30 days, and compared their observations to the generic predictions of the Bode & Kahn (1985) model. They concluded that their observations could come from the hot, post-shock gas – this model would predict that the IR source should be seen linearly expanding at a rate of about  $\sim 1$  milliarcseconds per day at a distance of 1.6 kpc. This distance estimate is based (most securely) on the expanding size of the radio emission observed in 1985 by Hjellming et al. (1986) and Taylor et al. (1989), assuming association with the forward shock (new radio data re-confirm the 1985 observations; Rupen et al. 2006; O’Brien et al. 2006). As will become clear, the distance to RS Oph is key to the interpretation of the IR interferometry data presented here.

Challenging this interpretation, Hachisu & Kato (2001, see also Kato 1991) have recently produced a comprehensive model for recurrent novae meant to

<sup>1</sup> monnier@umich.edu: University of Michigan Astronomy Department, 941 Dennison Bldg, Ann Arbor, MI 48109-1090, USA.

<sup>2</sup> Johns-Hopkins University

<sup>3</sup> Massachusetts Institute of Technology

<sup>4</sup> W. M. Keck Observatory

<sup>5</sup> Laboratoire d’Astrophysique de Grenoble

<sup>6</sup> Michelson Science Center (Caltech)

<sup>7</sup> Harvard-Smithsonian Center for Astrophysics

<sup>8</sup> Jet Propulsion Laboratory (Caltech)

<sup>9</sup> University of Massachusetts, Amherst

<sup>10</sup> California Institute of Technology

<sup>11</sup> Observatoire de Haute-Provence, France

<sup>12</sup> NASA – Goddard Space Flight Center

explain a wide range of the known novae properties, and makes a specific prediction for the origin of the near-infrared (NIR) continuum which is very different from Evans et al. (1988). Following onset of the thermonuclear runaway of the hydrogen shell around the WD, the shell expands to AU size, in effect turning the WD back into a red giant. The shell stably burns hydrogen for a few weeks, shrinking back to the size of white dwarf. According to this model, hot post-shock gas plays no role in forming the IR continuum. Furthermore, Hachisu & Kato (2001) prefer a much closer distance of 600 pc implying a binary separation of 2.9 milliarcseconds, easily detectable with current interferometers.

In this Letter, we report first-ever size measurements for RS Oph in the NIR using long-baseline interferometry. Our results are surprising, ruling out the favored expanding fireball model, raising doubts about the established distance to RS Oph, and motivating a new model for the NIR emission.

## 2. OBSERVATIONS

In the Letter we report on data from 3 different interferometers and a summary of observations can be found in Table 1. Here we briefly introduce each dataset.

Most of our data were obtained at the Infrared-Optical Telescope Array (IOTA, Traub et al. 2003) which has baselines between 5–38m. The IONIC3 combiner (Berger et al. 2003) was used to measure 3 visibilities ( $V^2$ ) and 1 closure phase (CP) simultaneously in the broadband H band filter ( $\lambda_0 = 1.65\mu\text{m}$ ,  $\Delta\lambda = 0.3\mu\text{m}$ ). Data analysis procedures have been documented in recent papers (Monnier et al. 2004, 2006). For the data here, we have adopted a calibration error  $\Delta V^2 = 0.05$  (relative error) on baselines AB and AC and  $\Delta V^2 = 0.10$  for baseline BC.

The Keck Interferometer (KI) was used to observe RS Oph on a single night 4 days after the burst (UT 2006 Feb 16) with a baseline of  $\sim 85\text{m}$ . Facility and instrument descriptions can be found in recent KI publications (Colavita et al. 2003; Colavita & Wizinowich 2003, 2000). The K-band ( $\lambda_0 = 2.18\mu\text{m}$ ,  $\Delta\lambda = 0.3\mu\text{m}$ ) data reported here were the by-product of a nulling observation (being prepared for a separate publication) and the calibration sequence is somewhat modified from the standard procedures; in particular, only one ratio measurement was made per integration.

Lastly, the Palomar Testbed Interferometer (PTI) observed RS Oph on 3 nights in 2006 April using the K-band detection system ( $\lambda_0 = 2.20\mu\text{m}$ ,  $\Delta\lambda = 0.4\mu\text{m}$ ) with a  $\sim 85\text{ m}$  baseline (oriented in NE direction). Detailed instrument and data analysis descriptions for PTI can be found in the literature (Colavita 1999; Colavita et al. 1999). Because of the inherent faintness of the source in April, coherent integration was used for analysis and a large calibration error of  $\Delta V^2 = 0.10$  (absolute, not relative) was added in quadrature with the internal error for model fitting in this paper.

We have split the data into 3 epochs – from 2006 Feb 16 to 23 (Days 4-11), from Feb 26 to Mar 13 (Days 14-29), and from April 2 to April 18 (Days 49-65). The UV-averaged visibility data for each epoch are presented in Figure 1 along with some Gaussian profiles for comparison. The IOTA closure phase results are shown in

Figure 2, also split into the three epochs.

All  $V^2$  and closure phase data are available from the authors; all data products are stored in the FITS-based, optical interferometry data exchange format (OI-FITS), recently described in Pauls et al. (2005).

## 3. ANALYSIS

Inspection of Figure 1 reveals that the visibility curve for RS Oph changes very little between day 4 and 65 (since the outburst). This is surprising since the H band brightness faded by a factor  $>10$  during this time. This result is discordant with the generic prediction of Evans et al. (1988) who interpreted the IR light curve in terms of the time-evolution of post-shock gas at  $10^5\text{K}$  moving at the speed of the contact discontinuity, 1400 km/s (following the earlier work by Bode & Kahn 1985). This model requires the IR emission to be seen expanding at a rate of  $\sim 1.0$  milliarcseconds per day (assuming a distance of 1.6kpc), an interpretation that is now ruled out.

Before discussing alternative models in §4, we wish to carry-out some model fitting to the interferometry data. Here we only consider two simple models – a circularly symmetric Gaussian and a binary star model. For all fits and calculation of reduced  $\chi^2$ , we have used the original data points before UV-averaging.

First, we fit a circularly symmetric Gaussian to each epoch of data, split by wavelength. Table 2 contains the best-fitting Full-Width at Half-Maxima (FWHM) and the reduced  $\chi^2$  for both the  $V^2$  and CP (the model CP is always zero for a Gaussian profile). The Gaussian model is a reasonable fit for the IOTA visibility alone, but clearly cannot fit the non-zero closure phase seen in March and April. Also, Figure 1 shows that no good fit was possible when combining IOTA with Keck and PTI data, indicating this model is too simplistic to explain the full range of baselines and/or the wavelength-dependence. Sizes derived from the longer-baseline K-band data are systematically smaller than those derived from shorter-baseline H-band data (IOTA).

As discussed in the Introduction, Hachisu & Kato (2001) suggest that the nova’s IR light curve might be due primarily to a rapid increase in brightness of the WD as it returns, albeit briefly, to a red giant phase. Motivated by this work, we realized that the IR emission might be due to the underlying RS Oph binary itself, and this might explain the general puzzling features of our data: non-expanding emission size, the inadequacies of the Gaussian fit, and the non-zero closure phase.

In order to test this idea, we fit binary models to the data for each of the three time periods, treating the brightness ratio as independent of wavelength in order to fit the H and K band data together. The IOTA, KI, and PTI complement each other in Fourier coverage and an exhaustive grid search of separations less than 10 mas found unique binary star solutions<sup>13</sup>. Table 2 contains the best-fit binary models for the three time periods, including the reduced  $\chi^2$  ( $V^2$ , CP). All three epochs are reasonably fit by a similar binary model. The only parameter that changed significantly between the epochs

<sup>13</sup> Our limited (u,v) coverage admit some unlikely additional solutions with larger binary separations which will be discussed fully in a future modeling paper.

was the brightness ratio. The closure phase predictions for the binary models are plotted along with the closure phase data in Figure 2. We note that our uv-plane is missing coverage in the NW direction and thus elongated structure in this direction would be observed foreshortened.

#### 4. DISCUSSION

Because the expanding fireball model fails to explain the nearly static size scale of the IR emission, we now seek suitable alternative emission mechanisms for the time-variable IR emission from the recurrent nova RS Oph. We have pursued the reborn red giant (thermonuclear runaway) model of Hachisu & Kato (2001) and found that indeed our 3-interferometer combined dataset can be explained by a simple binary model with separation of  $\sim 3.2$  milliarcsecond, PA  $30^\circ$  E of N, and a brightness ratio varying from 2.5:1 to 5:1. Next we subject the binary hypothesis to further scrutiny.

##### 4.1. Binary Interpretation of Near-Interferometry Data

Based on single-line radial velocity data, Fekel et al. (2000) finds the RS Oph binary orbit to be roughly circular with a period of  $455.72 \pm 0.83$  days and mass function  $f = 0.221 \pm 0.038 M_\odot$ . RGs in symbiotic systems are typically  $1-3 M_\odot$  (Dobrzycka & Kenyon 1994) and we expect recurrent novae to contain a Chandrasekhar mass WD ( $1.4 M_\odot$ ); these facts combined with the known mass function rule out RG masses greater than  $2 M_\odot$ . Assuming the RS Oph system mass to be  $2.4-3.4 M_\odot$ , we find the component separation to be  $1.55-1.74$  AU, or (unprojected)  $2.59-2.90$  mas at the 600 pc distance preferred by Hachisu & Kato (2001) – only slightly smaller than our observed separation of 3.2 mas.

Since the RS Oph outburst took place only 1 month before maximum redshifted velocity (Fekel et al. 2000), our measured binary parameters represent the true orbital semi-major axis and orbital  $\Omega$  for RS Oph with only weak  $\sin(i)$  dependencies. Thus, a small reduction in the distance estimate (540 pc) brings the interferometer binary model in agreement with expectations from Kepler’s laws.

The binary model fits (Table 2) show evidence for a change in the brightness ratio over time. While the Hachisu & Kato (2001) theory predicts a time-changing brightness ratio, it is beyond the scope of this paper to test the compatibility with the observed IR light curves due to complications from the role of the irradiated RG photosphere and the presence of a post-outburst WD accretion disk.

##### 4.2. Circumbinary Reservoir of Hot Gas

The distance estimate of  $\lesssim 540$  pc derived in the last section stands in strong contrast to estimates more commonly adopted in the literature. The most significant constraints on distance are set by resolved radio observations of the previous and current burst (Hjellming et al. 1986; Taylor et al. 1989; Rupen et al. 2006; O’Brien et al. 2006). By assuming that observed radio proper motions (on the sky) can be ascribed to the fast-moving ejecta or forward shock, workers consistently derive a distance of  $\sim 1600$  pc. Similarly large distances were found considering interstellar UV absorption

lines (Snijders 1987) and HI absorption measurements (Hjellming et al. 1986).

Given the strength of the evidence, we now consider the implications of the  $d = 1600$  pc distance. This distance would rule out the binary interpretation of the near-IR interferometry data laid out in §4.1, given existing binary constraints. Instead, we hypothesize that the IR emission arises from a quasi-stationary<sup>14</sup> hot gas reservoir that contributes a combination of emission lines and free-free/bound-free emission in the NIR bands. The observed FWHM of  $\sim 3$  mas is  $\sim 5$  AU at 1600 pc, about  $3\times$  the expected RS Oph binary separation. This size is reasonable for a circumbinary disk or reservoir of hot gas, perhaps kept ionized by the outward moving blastwave or soft X-ray luminosity from the WD itself following outburst. This gas reservoir might be analogous to the “fallback disk” inferred to form after some supernovae (e.g., Wang et al. 2006).

Clearly, the hypothesized gas reservoir must be elongated and somewhat off-center with respect to the central source in order to fit the combined IOTA, KI, and PTI interferometry data, especially the non-zero closure phases. It is beyond the scope of this Letter to investigate the details here, and we defer development of this model to a future paper.

##### 4.3. Future Work

While we have ruled out the important class of expanding fireball models for explaining the IR emission from the recurrent nova RS Oph, more work lies ahead to test the other emission mechanisms discussed in this Letter. A future study will attempt to synthesize a *self-consistent* model that can explain the time evolution of the IR spectrum, NIR and mid-IR interferometer data, and multi-wavelength light curves at the same time<sup>15</sup>.

If the close distance  $d \lesssim 540$  pc is confirmed, we have a spectacular opportunity to study in detail a likely SNIa progenitor and to learn about unexpected shock physics controlling the non-thermal radio emission. Alternatively, the further distance  $d \sim 1600$  pc suggests we have discovered a significant and new component to the RS Oph Nova remnant and future work will characterize the hot circumbinary gas reservoir for the first time.

We thank J. Sokoloski for productive discussion and comments on the manuscript and members of the AAVSO IR group for early-time photometry. R.K.B. acknowledges NASA and Dr. Bill Danchi for supporting this work. We acknowledge SAO, NSF, NASA, CNRS/CNES (France), LAOG, IMEP, and LETI for support of IOTA and IONIC3. KI was developed and is operated by JPL, MSC and WMKO with funding from NASA. PTI was developed by JPL and is operated by the MSC on behalf of the PTI collaboration. This research made use of SIMBAD, ADS, MSC resources, CHARM2, and the 2MASS catalog. Some data presented herein

<sup>14</sup> It is possible to fit our data with an expanding wind or jet component, but this requires fine-tuning the relative proportions of multiple components and/or a very asymmetric jetlike emission oriented perpendicular to our long (northeast) baselines. These possibilities will be investigated more thoroughly in future work with an extended dataset.

<sup>15</sup> Note added: future modeling should address the asymmetric radio nebula and jet observed by O’Brien et al. (2006).

TABLE 1  
OBSERVING LOG FOR RS OPH

Days since 2006 Feb 12	Date (UT)	Interferometer (Configuration)	Wavelength ( $\mu\text{m}$ )
4	2006Feb16	IOTA (A20B15C00) <sup>a</sup>	1.65
4	2006Feb16	Keck <sup>b</sup>	2.18
11	2006Feb23	IOTA (A20B15C00)	1.65
14	2006Feb26	IOTA (A35B15C10)	1.65
19	2006Mar03	IOTA (A35B15C10)	1.65
20	2006Mar04	IOTA (A35B15C10)	1.65
22	2006Mar06	IOTA (A35B15C10)	1.65
25	2006Mar09	IOTA (A35B15C10)	1.65
29	2006Mar13	IOTA (A35B15C10)	1.65
49	2006Apr02	Palomar Testbed (NW) <sup>c</sup>	2.20
59	2006Apr12	IOTA (A35B15C10)	1.65
60	2006Apr13	IOTA (A35B15C10)	1.65
63	2006Apr16	Palomar Testbed (NW)	2.20
65	2006Apr18	Palomar Testbed (NW)	2.20

NOTE. — Uniform Disk (UD) diameters of calibrators were generally estimated using *getCal*, an SED-fitting routine maintained and distributed by the Michelson Science Center (<http://msc.caltech.edu>).

<sup>a</sup>IOTA used the following calibrators: HD 152601 ( $1.6\pm0.4$  mas), HD 164064 ( $1.6\pm0.5$  mas),  $\chi$  UMa ( $3.24\pm0.04$  mas; Bordé et al. 2002),  $\rho$  Boo ( $3.72\pm0.12$  mas; van Belle et al. 1999), HD 143033 ( $1.9\pm1.5$  mas), HD 156826 ( $0.6\pm0.2$  mas), HD 157262 ( $1.5\pm0.5$  mas)

<sup>b</sup>KI used the following calibrators:  $\chi$  UMa ( $3.35\pm0.17$  mas; Cohen et al. 1999),  $\rho$  Boo ( $3.92\pm0.19$  mas; Cohen et al. 1999).

<sup>c</sup>PTI used the following calibrators: HD 164064 ( $1.6\pm0.5$  mas), HD 161868 ( $0.7\pm0.1$  mas).

were obtained at the Keck Observatory, operated by a

scientific partnership among Caltech, UC, and NASA.

#### REFERENCES

- Berger, J.-P., *et al.* 2003, in *Interferometry for Optical Astronomy II*. Edited by Wesley A. Traub. Proceedings of the SPIE, Volume 4838, 1099–1106
- Bode, M. F. & Kahn, F. D. 1985, *MNRAS*, 217, 205
- Bordé, P., Coudé du Foresto, V., Chagnon, G., & Perrin, G. 2002, *A&A*, 393, 183
- Cohen, M., Walker, R. G., Carter, B., Hammersley, P., Kidger, M., & Noguchi, K. 1999, *AJ*, 117, 1864
- Colavita, M., *et al.* 2003, *ApJ*, 592, L83
- Colavita, M. M. 1999, *PASP*, 111, 111
- Colavita, M. M., *et al.* 1999, *ApJ*, 510, 505
- Colavita, M. M. & Wizinowich, P. L. 2000, in *Proc. SPIE Vol. 4006*, p. 310–320, *Interferometry in Optical Astronomy*, Pierre J. Lena; Andreas Quirrenbach; Eds., 310–320
- Colavita, M. M. & Wizinowich, P. L. 2003, in *Interferometry for Optical Astronomy II*. Edited by Wesley A. Traub. Proceedings of the SPIE, Volume 4838., 79–88
- Dobrzycka, D. & Kenyon, S. J. 1994, *AJ*, 108, 2259
- Evans, A., Callus, C. M., Albinson, J. S., Whitelock, P. A., Glass, I. S., Carter, B., & Roberts, G. 1988, *MNRAS*, 234, 755
- Fekel, F. C., Joyce, R. R., Hinkle, K. H., & Skrutskie, M. F. 2000, *AJ*, 119, 1375
- Gehrz, R. D. 1988, *ARA&A*, 26, 377
- Hachisu, I. & Kato, M. 2001, *ApJ*, 558, 323
- Hjellming, R. M., van Gorkom, J. H., Seaquist, E. R., Taylor, A. R., Padin, S., Davis, R. J., & Bode, M. F. 1986, *ApJ*, 305, L71
- Kato, M. 1991, *ApJ*, 369, 471
- Lane, B. F., Retter, A., Eisner, J. A., Thompson, R. R., & Mutterspaugh, M. W. 2005, *American Astronomical Society Meeting Abstracts*, 207,
- Monnier, J. D., Berger, J.-P., & others. 2006, *ApJ*, in press
- Monnier, J. D., *et al.* 2004, *ApJ*, 602, L57
- Narumi, H., Hirosawa, K., Kanai, K., Renz, W., Pereira, A., Nakano, S., Nakamura, Y., & Pojmanski, G. 2006, *IAU Circ.*, 8671, 2
- O’Brien, T. J., *et al.* 2006, *Nature*, in press
- Pauls, T. A., Young, J. S., Cotton, W. D., & Monnier, J. D. 2005, *PASP*, 117, 1255
- Rupen, Mioduszewski, & Sokoloski. 2006, in preparation
- Snijders, M. A. J. 1987, in *RS Ophiuchi (1985) and the Recurrent Nova Phenomenon*, ed. M. F. Bode, 51–+
- Taylor, A. R., Davis, R. J., Porcas, R. W., & Bode, M. F. 1989, *MNRAS*, 237, 81
- Traub, W. A., *et al.* 2003, in *Interferometry for Optical Astronomy II*. Edited by Wesley A. Traub. Proceedings of the SPIE, Volume 4838, pp. 45–52 (2003)., 45–52
- van Belle, G. T., *et al.* 1999, *AJ*, 117, 521
- Wang, Z., Chakrabarty, D., & Kaplan, D. L. 2006, *Nature*, 440, 772
- Warner, B. 1976, in *IAU Symp. 73: Structure and Evolution of Close Binary Systems*, ed. P. Eggleton, S. Mitton, & J. Whelan, 85–+

TABLE 2  
RS OPH MODEL FITTING RESULTS

Model Parameter	2006 Feb 16 – 23 (Days 4 – 11)	2006 Feb 26 – Mar 13 (Days 14 – 29)	2006 Apr 02 – 18 (Days 49 – 65)
Gaussian Profile (fitting only to $1.65\mu\text{m}$ )			
FWHM (milliarcseconds)	$3.30 \pm 0.09$	$3.47 \pm 0.03$	$2.87 \pm 0.07$
Reduced $\chi^2$ ( $V^2$ )	0.6	1.3	1.1
Reduced $\chi^2$ (CP)	1.3	3.6	5.4
Gaussian Profile (fitting only to $2.2\mu\text{m}$ )			
FWHM (milliarcseconds)	$2.56 \pm 0.24$	N/A	$2.00 \pm 0.09$
Reduced $\chi^2$ ( $V^2$ )	0 <sup>a</sup>	N/A	1.0
Reduced $\chi^2$ (CP)	N/A	N/A	N/A
Binary Model <sup>b</sup> (fitting to $1.65\mu\text{m}$ & $2.2\mu\text{m}$ data)			
Separation (milliarcseconds)	$3.13 \pm 0.12$	$3.23 \pm 0.13$	$3.48 \pm 0.23$
Position Angle <sup>c</sup> (degs E of N)	$36 \pm 10^d$	$45 \pm 5$	$27 \pm 5$
Brightness Ratio <sup>c</sup>	$0.42 \pm 0.06$	$0.40 \pm 0.06$	$0.21 \pm 0.03$
Reduced $\chi^2$ ( $V^2$ )	0.6	1.7	1.2
Reduced $\chi^2$ (CP)	1.1	1.3	0.5

<sup>a</sup>Only one Keck data point available for fitting.

<sup>b</sup>Brightness ratio assumed the same for H and K bands. Individual components have adopted UD diameters of 0.5 mas. Here, we considered only solutions with separations  $< 10$  mas.

<sup>c</sup>Fainter component with respect to brighter component.

<sup>d</sup> $180^\circ$  ambiguity since closure phase data are indistinguishable from zero for this epoch.

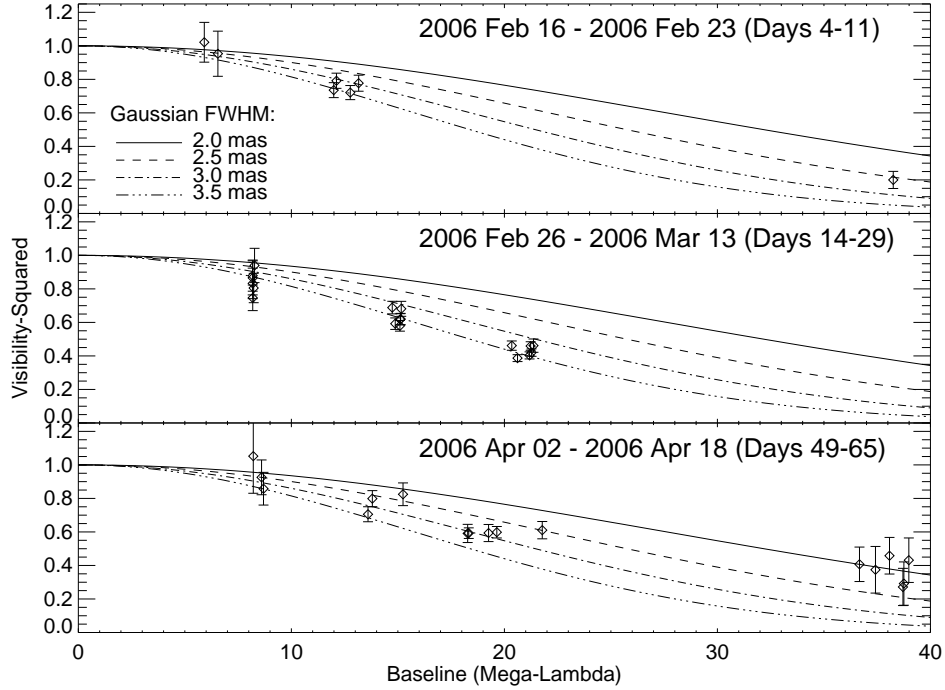


FIG. 1.— This figure shows UV-averaged visibility data for RS Oph split into three different time periods. The x-axis shows the spatial frequency (projected baseline in units of wavelength) while the y-axis shows the visibility-squared. Four curves representing Gaussian profiles are also included to show the characteristic size and to allow intercomparison of data in the different panels. For reference, all data shortward of  $25\text{M}\lambda$  derive from IOTA, while longer baseline data come from KI (Epoch I) and PTI (Epoch III).

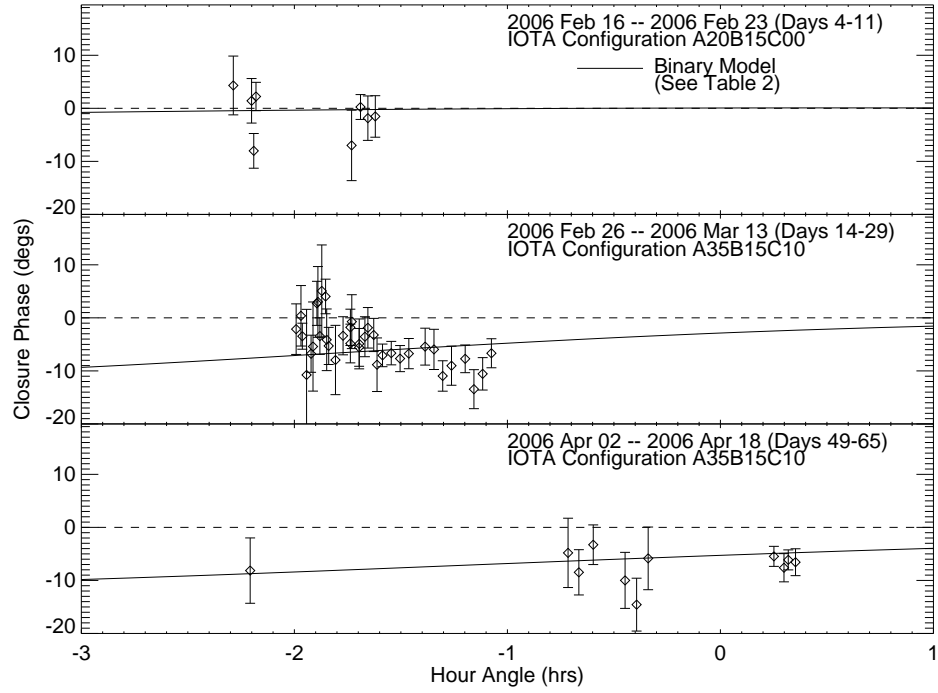


FIG. 2.— This figure shows IOTA closure phase data for RS Oph split into three different time periods. The x-axis shows the hour angle of the observations while the y-axis shows the observed closure phase. The solid line shows expected closure phase for the binary model parameters found in Table 2 and discussed in §3. The dashed line shows the expected closure phase signal for the symmetric Gaussian model.